



The *Suzaku* view of highly ionized outflows in AGN – II. Location, energetics and scalings with bolometric luminosity

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ABSTRACT

Ongoing studies with *XMM–Newton* have shown that powerful accretion disc winds, as revealed through highly ionized Fe K-shell absorption at $E \geq 6.7$ keV, are present in a significant fraction of active galactic nuclei (AGNs) in the local Universe (Tombesi et al. 2010a). In Gofford et al., we analysed a sample of 51 *Suzaku*-observed AGNs and independently detected Fe K absorption in ~ 40 per cent of the sample, and we measured the properties of the absorbing gas. In this work, we build upon these results to consider the properties of the associated wind. On average, the fast winds ($v_w > 0.01c$) are located $\langle r \rangle \sim 10^{15-18}$ cm (typically $\sim 10^{2-4} r_s$) from their black hole, their mass outflow rates are of the order of $\langle \dot{M}_w \rangle \sim 0.01-1 M_\odot \text{ yr}^{-1}$ or $\sim (0.01-1) \dot{M}_{\text{Edd}}$ and kinetic power is constrained to $\langle L_w \rangle \sim 10^{43-45} \text{ erg s}^{-1}$, equivalent to $\sim (0.1-10 \text{ per cent}) L_{\text{Edd}}$. We find a fundamental correlation between the source bolometric luminosity and the wind velocity, with $v_w \propto L_{\text{bol}}^\alpha$ and $\alpha = 0.4_{-0.2}^{+0.3}$ (90 per cent confidence), which indicates that more luminous AGN tend to harbour faster Fe K winds. The mass outflow rate \dot{M}_w , kinetic power L_w and momentum flux \dot{p}_w of the winds are also consequently correlated with L_{bol} , such that more massive and more energetic winds are present in more luminous AGN. We investigate these properties in the framework of a continuum-driven wind, showing that the observed relationships are broadly consistent with a wind being accelerated by continuum-scattering. We find that, globally, a significant fraction (~ 85 per cent) of the sample can plausibly exceed the $L_w/L_{\text{bol}} \sim 0.5$ per cent threshold thought necessary for feedback, while 45 per cent may also exceed the less conservative ~ 5 per cent of L_{bol} threshold as well. This suggests that the winds may be energetically significant for AGN–host-galaxy feedback processes.

Key words: line: identification – galaxies: active – galaxies: nuclei – X-rays: galaxies.

1 INTRODUCTION

Outflows of photoionized gas are ubiquitous in active galactic nuclei (AGNs). Recent surveys with *XMM–Newton* and *Suzaku* by Tombesi et al. (2010b) and Gofford et al. (2013, hereafter Paper I) have shown that Fe K-shell outflows are observed in a significant fraction ($\sim 40-50$ per cent) of active galaxies. These potentially massive outflows ($N_H \sim 10^{23} \text{ cm}^{-2}$), which are identified through blueshifted resonant absorption from Fe XXV He α

($E_{\text{rest}} = 6.7$ keV) and Fe XXVI Ly α ($E_{\text{rest}} = 6.97$ keV), can possess outflow velocities (v_w) ranging from a few thousand km s $^{-1}$ up to mildly relativistic values ($v_w \sim 0.1-0.3c$; Chartas et al. 2002; Pounds et al. 2003; Reeves et al. 2009; Tombesi et al. 2010a, 2015; Paper I) which indicates a substantial mass transport into the host galaxy. The large wind velocity – which indicates an origin directly associated with the accretion disc, hence leading them to be dubbed ‘disc-winds’ – implies that the ensuing outflow may be energetically significant in terms of feedback (i.e. $L/L_{\text{bol}} \gtrsim 0.5-5$ per cent; Di Matteo, Springel & Hernquist 2005; Hopkins & Elvis 2010). The winds are observed in both radio-quiet (Tombesi et al. 2010a; Paper I) and radio-loud (Paper I; Tombesi et al. 2014) AGNs which

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suggests that they may also be an important addition to the traditional AGN unified model (e.g. Antonucci 1993; Urry & Padovani 1995).

In Paper I, we performed a systematic search for Fe K absorption in 51 *Suzaku*-observed AGN which were heterogeneously selected from the HEASARC data archive¹ (see Paper I for selection criteria). Our main results from that paper were as follows: (i) blueshifted Fe K-shell absorption is present at the $\geq 2\sigma$ level in 20/51 (~ 40 per cent) of the selected AGN (in 28/73 individual observations); (ii) the absorbing gas is typified by mean values of $\log(N_{\text{H}}/\text{cm}^{-2}) \sim 23$ and $\log(\xi/\text{erg cm s}^{-1}) \sim 4.5$, while the outflow velocity spans a continuous range between $v_{\text{w}} < 1500 \text{ km s}^{-1}$ and $\lesssim 100\,000 \text{ km s}^{-1}$, respectively, with a mean and median velocity of $0.1c$ and $0.056c$. Ultimately, these results are all consistent with those found by Tombesi et al. (2010a) on the basis of their *XMM-Newton* sample.

In this work, we build upon our previous results to assess the properties of the disc-wind itself. The remainder of this paper is laid out as follows: in Section 2, we outline the equations through which assess the wind properties, in Section 3.2, we use the inferred values to perform a qualitative assessment of the wind properties, before going on to conduct a quantitative correlation analysis of how the various wind properties scale with the AGN bolometric luminosity in Section 3.3. In Section 4, we then compare our results to those already present in the literature, discuss the wind acceleration mechanism which may be responsible for the observed outflows and finally assess whether the observed wind is likely to be energetically significant in terms of feedback. We summarize our results and present our overall conclusions in Section 5.

2 WIND PARAMETERS

2.1 Projected distance

The maximum column density N_{H} of gas along the line of sight (LOS) is given by $N_{\text{H}} = \int_{r_{\text{max}}}^{\infty} n(r) dr$, where $n(r)$ is the average gas number density and r_{max} is the maximum observed distance of the absorber. This, combined with the definition of the ionization parameter $\xi = L_{\text{ion}}/nr^2$ (Tarter, Tucker & Salpeter 1969), yields the maximum distance that the absorber can be located from the ionizing source given its observed column density and ionization state,

$$r_{\text{max}} = \frac{L_{\text{ion}}}{\xi N_{\text{H}}}, \quad (1)$$

where L_{ion} is the source ionizing luminosity integrated between 1 and 1000 Rydberg. Conversely, a lower limit on r can be inferred by considering the escape radius of the gas given its observed velocity. For a simple Keplerian disc orbiting a black hole (BH), the escape velocity at distance r is $v_{\text{esc}} = \sqrt{2GM_{\text{BH}}/r}$. In the limit that $v_{\text{w}} = v_{\text{esc}}$, i.e. assuming that the measured outflow velocity along the LOS is equal to the escape velocity at observed radius r , we can set a lower limit on the location of the wind,

$$r_{\text{min}} = \frac{2GM_{\text{BH}}}{v_{\text{w}}^2}. \quad (2)$$

These relations hence allow upper/lower limits to be placed on the wind location given the measured parameters of the absorbing gas, albeit with large uncertainties.

2.2 Mass outflow rate

The mass outflow rate of the wind, $dM/dt = \dot{M}_{\text{w}}$, is a crucial parameter and is the main means through which the overall flow energetics are assessed. The mass outflow rate for an arbitrary wind is given by $\dot{M}_{\text{w}} = A(r)\rho(r)v(r)$, where $\rho(r)$ and $v(r)$ are the density and velocity profile of the wind and $A(r)$ is a factor which accounts for the geometry of the system. Assuming that the flow has constant terminal velocity $v(r) = v_{\text{w}}$ and that the absorbing gas has cosmic elemental abundances (i.e. ~ 75 per cent of its mass by hydrogen and ~ 25 per cent by helium), $\rho(r) \simeq 1.2m_{\text{p}}n(r)$, where m_{p} is the proton mass and $n(r)$ is the electron number density of the plasma. For a thin spherically symmetric isotropic wind $A(r) = \Omega br^2$, where the product $\Omega b \leq 1$ is known as the *global filling factor* and accounts for both the solid angle occupied by the flow (Ω) and how much of the flow volume is filled by gas (b). Thus, $\dot{M}_{\text{w}} \sim \Omega br^2 m_{\text{p}} n(r) v_{\text{w}}$ where we have neglected the constant factor of order unity.

The value b is extremely difficult to determine because it depends on the ionization and clumpiness of the gas. At low-intermediate ionization states the flow is likely to be clumpy/filamentary, while at high-ionization states it can be considered largely smooth and of low density because the vast majority of elements are stripped of electrons. In the clumpy case, the column density of the wind can be given by $N_{\text{H}} \sim bn(r)\delta r$, with $b < 1$ implicitly allowing for inhomogeneities in the flow. Alternatively, at high-ionization states, $b \simeq 1$ and $r \rightarrow r_{\text{max}}$. Substituting for r_{min} and r_{max} then leads to algebraic upper and lower limits on the mass outflow rate:

$$\dot{M}_{\text{w}}^{\text{max}} \sim \Omega m_{\text{p}} L_{\text{ion}} \xi^{-1} v_{\text{w}}, \quad (3a)$$

$$\dot{M}_{\text{w}}^{\text{min}} \sim 2\Omega G M_{\text{BH}} m_{\text{p}} N_{\text{H}} v_{\text{w}}^{-1}. \quad (3b)$$

The major remaining uncertainty is the opening angle of the system Ω . Here, we adopt the average opening angle $\Omega = 1.6\pi$ as inferred from the global detection fraction of Fe K winds reported in the literature ($f \simeq 40$ per cent; Tombesi et al. 2010a, 2014; Paper I). The assumption of a uniform wind geometry is clearly an oversimplification of real systems; in reality, numerous factors, such as source luminosity, gas density, etc., will contribute to the shaping of an X-ray disc-wind and therefore the wind geometry is going to differ on an object-by-object basis. Even so, we note that the average opening angle of $\Omega = 1.6\pi$ is comparable to the wind opening measured from the P-Cygni-like Fe K profile in PDS 456 (Ω is resolved to be $> 2\pi$; see Nardini et al. 2015). This suggests that it is a good approximation to the geometry of real disc-winds.

2.3 Kinetic power

Provided that the wind has already reached a steady terminal velocity by the point at which it is observed, the mechanical power imparted by expelling mass at a rate \dot{M}_{w} with velocity v_{w} is simply equal to its kinetic energy: $L_{\text{w}} = \dot{M}_{\text{w}} v_{\text{w}}^2 / 2$. We determine the range of likely kinetic power by substituting for $\dot{M}_{\text{w}} = (\dot{M}_{\text{w}}^{\text{max}}, \dot{M}_{\text{w}}^{\text{min}})$, leading to

$$L_{\text{w}}^{\text{max}} \sim \Omega m_{\text{p}} L_{\text{ion}} \xi^{-1} v_{\text{w}}^3 \quad (4a)$$

$$L_{\text{w}}^{\text{min}} \sim G\Omega m_{\text{p}} M_{\text{BH}} N_{\text{H}} v_{\text{w}}. \quad (4b)$$

Similarly, the rate at which the outflow transports momentum into the environment of the host galaxy is given by $dp/dt \equiv$

¹ Accessible at <http://heasarc.gsfc.nasa.gov/docs/archive.html>.

Table 1. Summary of AGN and wind properties. Columns: (1) SMBH mass, in units of M_\odot ; (2) absorption-corrected 1–1000 Rydberg ionizing luminosity (in units of erg s^{-1}), as determined from the best-fitting continuum model outlined in Paper I; (3) measured wind column density, in units of cm^{-2} ; (4) wind ionization parameter, in units of erg cm s^{-1} ; (5) inferred wind outflow velocity, as measured in Paper I.

| Source | $\log M_{\text{BH}}$ (1) | $\log L_{\text{ion}}$ (2) | $\log N_{\text{H}}$ (3) | $\log \xi$ (4) | v_w/c (5) |
|----------------------------------|-----------------------------|------------------------------|----------------------------|------------------------|---------------------------|
| 3C 111 ^{1,2,*} | 8.1 ± 0.5 | 44.7 | $23.0^{+0.4}_{-0.8}$ | $4.45^{+0.40}_{-0.46}$ | $0.072^{+0.041}_{-0.038}$ |
| 3C 390.3 ^{3,4} | $8.8^{+0.2}_{-0.6}$ | 44.7 | >23.7 | >5.46 | 0.145 ± 0.007 |
| 4C +74.26 ⁵ | 9.6^\dagger | 47.0 | >21.8 | 4.06 ± 0.45 | 0.185 ± 0.026 |
| APM 08279 ^{6,a} | 10.0^\dagger | 45.4 | 23.0 ± 0.1 | $3.51^{+0.33}_{-0.18}$ | $0.285^{+0.165}_{-0.158}$ |
| CBS 126 ⁷ | 7.8 ± 0.1 | 44.2 | >23.7 | $4.77^{+0.26}_{-0.17}$ | 0.012 ± 0.006 |
| ESO 103-G035 ^{8,*} | 7.4 ± 0.1 | 43.9 | >21.9 | 4.36 ± 1.19 | 0.056 ± 0.025 |
| MCG -6-30-15 ⁹ | 6.7 ± 0.2 | 43.2 | 22.2 ± 0.1 | 3.64 ± 0.06 | 0.007 ± 0.002 |
| MR 2251-178 ¹⁰ | 8.7 ± 0.1 | 45.2 | 21.5 ± 0.2 | 3.26 ± 0.12 | 0.137 ± 0.008 |
| Mrk 279 ⁴ | 7.5 ± 0.1 | 43.4 | 23.4 ± 0.3 | $4.42^{+0.15}_{-0.27}$ | 0.220 ± 0.006 |
| Mrk 766 ¹¹ | $6.2^{+0.3}_{-0.6}$ | 43.2 | $22.7^{+0.2}_{-0.3}$ | $3.86^{+0.37}_{-0.25}$ | $0.039^{+0.030}_{-0.026}$ |
| NGC 1365 ^{12,*} | 7.6 ± 0.7 | 42.8 | $23.7^{+0.2}_{-0.5}$ | 3.88 ± 0.07 | <0.014 |
| NGC 3227 ³ | 7.6 ± 0.2 | 42.5 | $22.7^{+0.2}_{-0.3}$ | $3.88^{+0.12}_{-0.16}$ | <0.008 |
| NGC 3516 ³ | 7.6 ± 0.1 | 43.6 | 22.6 ± 0.2 | 3.84 ± 0.11 | 0.004 ± 0.002 |
| NGC 3783 ³ | 7.5 ± 0.1 | 43.6 | 21.8 ± 0.2 | $3.48^{+0.15}_{-0.07}$ | <0.008 |
| NGC 4051 ³ | 6.3 ± 0.2 | 42.5 | 22.8 ± 0.1 | $3.50^{+0.53}_{-0.50}$ | $0.018^{+0.004}_{-0.005}$ |
| NGC 4151 ³ | 7.1 ± 0.1 | 42.9 | >21.7 | 3.69 ± 0.64 | 0.055 ± 0.023 |
| NGC 4395 ¹³ | 4.7 ± 0.2 | 40.7 | 22.8 ± 0.3 | 3.92 ± 0.16 | <0.001 |
| NGC 5506 ^{10, 14, 15,*} | 7.3 ± 0.7 | 43.7 | 23.2 ± 0.3 | $5.04^{+0.29}_{-0.17}$ | 0.246 ± 0.006 |
| PDS 456 ^{16,*} | 9.4 ± 0.3 | 45.3 | 23.0 ± 0.1 | $4.06^{+0.28}_{-0.15}$ | 0.273 ± 0.028 |
| SW J2127 ^{17,b} | 7.2^\dagger | 43.7 | 22.8 ± 0.3 | $4.16^{+0.29}_{-0.13}$ | 0.231 ± 0.006 |

References: ¹Chatterjee et al. (2011); ²Tombesi et al. (2012b); ³Peterson et al. (2004); ⁴Bentz et al. (2009a); ⁵Woo & Urry (2002); ⁶Saez, Chartas & Brandt (2009); ⁷Grupe et al. (2004); ⁸Czerny et al. (2001); ⁹McHardy et al. (2005); ¹⁰Khorunzhev et al. (2012); ¹¹Bentz et al. (2009b); ¹²Risaliti et al. (2007); ¹³Edri et al. (2012); ¹⁴Papadakis (2004); ¹⁵Nikolajuk, Czerny & Gurynowicz (2009); ¹⁶Reeves et al. (2009); ¹⁷Malizia et al. (2008).

Notes. * M_{BH} is taken as the mean of extreme values found in the literature; † errors on M_{BH} not present in the literature; ^afull designation APM 08279+5255; ^bfull designation SWIFT J2127.4+5654. Where a source has an Fe K wind detected in more than one observation, the mean N_{H} , ξ and v_w values are reported here; see Paper I for further details.

$\dot{p}_{\text{out}} = \dot{M}_w v_w$. Substituting again for \dot{M}_w leads to a plausible range between

$$\dot{p}_w^{\text{max}} \sim \Omega m_p L_{\text{ion}} \xi^{-1} v_w^2 \quad (5a)$$

$$\dot{p}_w^{\text{min}} \sim G \Omega m_p M_{\text{BH}} N_{\text{H}}. \quad (5b)$$

2.4 Other parameters

2.4.1 BH masses

Our method of estimating the inner radius of the outflow, r_{min} , is proportional to the escape radius of the BH, r_{esc} , which in turn is proportional to the mass of the central BH, M_{BH} . By extension, this also means that the other lower limiting quantities in equations (2), (3b), (4b) and (5b) are also proportional to M_{BH} . Estimates for M_{BH} gathered from the literature are collated in Table 1. We obtain most of the M_{BH} values through the numerous reverberation mapping studies available in the literature (i.e. Peterson et al. 2004; Bentz et al. 2009a,b; Edri et al. 2012) which tend to offer relatively tight con-

straints on the mass of the central object; in the cases where an AGN had been subject to reverberation mapping multiple times, we adopt the most recent estimate on M_{BH} only. There are several empirically determined mass estimates for the BHs in 3C 111, ESO 103-G035, NGC 1365, NGC 5506 and PDS 456. For these AGNs, we report the mean value, with the associated errors taken as half of the range between the minimum and maximum values to account for the uncertainty in the individual estimates on M_{BH} . For CBS 126, where we were unable to find any robust mass estimates in the literature, we infer M_{BH} from the BLR (broad-line region) line width and luminosity scaling relation of Kaspi et al. (2000) using the appropriate spectral values listed in Grupe et al. (2004). For 3C 390.3, we use the range of values inferred by Dietrich et al. (2012). In all other sources, we take M_{BH} directly from the reference listed in the table footnote. The sample encompasses almost six orders of magnitude in BH mass, with $\log M_{\text{BH}}$ ranging from ~ 4.7 in the dwarf-Seyfert galaxy NGC 4395 (Edri et al. 2012), all the way up to an estimated ~ 10 in the high-luminosity BAL quasar APM 08279+5255 (Saez et al. 2009).

Table 2. Summary of inferred wind parameters.

| Source | $\log r_{\min}$ (cm) | $\log r_{\max}$ (cm) | $\log \dot{M}_{\text{w}}^{\min}$ (g s^{-1}) | $\log \dot{M}_{\text{w}}^{\max}$ (g s^{-1}) | $\log L_{\text{w}}^{\min}$ (erg s^{-1}) | $\log L_{\text{w}}^{\max}$ (erg s^{-1}) |
|--------------|-------------------------|-------------------------|---|---|---|---|
| 3C 111 | 15.9 ± 0.1 | 17.3 ± 0.6 | 24.9 ± 0.1 | 26.4 ± 0.6 | 43.3 ± 0.1 | 44.7 ± 0.6 |
| 3C 390.3 | 15.7 ± 0.1 | <16.7 | >25.5 | <26.6 | >44.5 | <45.6 |
| 4C +74.26 | 16.5 ± 0.1 | <20.0 | >24.8 | 27.9 ± 0.5 | >44.0 | $47.1^{+0.5}_{-0.7}$ |
| APM 08279 | 16.6 ± 0.1 | 19.8 ± 0.3 | 26.3 ± 0.1 | 29.6 ± 0.3 | 45.9 ± 0.1 | 49.1 ± 0.3 |
| CBS 126 | — | <15.9 | — | 24.8 ± 0.3 | — | 41.6 ± 0.2 |
| ESO 103-G035 | $15.4^{+0.5}_{-0.3}$ | <18.8 | >23.0 | 25.6 ± 1.2 | >41.1 | 43.7 ± 1.2 |
| MCG -6-30-15 | 16.5 ± 0.3 | 17.4 ± 0.1 | 23.8 ± 0.3 | 24.7 ± 0.1 | 40.1 ± 0.3 | 41.0 ± 0.1 |
| MR 2251-178 | 15.9 ± 0.1 | 20.4 ± 0.3 | 24.3 ± 0.1 | 27.8 ± 0.7 | 43.3 ± 0.1 | 46.7 ± 0.7 |
| Mrk 279 | 14.3 ± 0.1 | 15.6 ± 0.5 | 24.3 ± 0.1 | 25.6 ± 0.5 | 43.6 ± 0.1 | 44.9 ± 0.5 |
| Mrk 766 | 14.5 ± 0.1 | 16.7 ± 0.3 | 23.1 ± 0.1 | 25.2 ± 0.3 | 40.9 ± 0.1 | 43.1 ± 0.3 |
| NGC 1365 | — | 15.2 ± 0.1 | — | 24.1 ± 0.1 | — | 40.5 ± 0.1 |
| NGC 3227 | — | 16.0 ± 0.4 | — | 23.6 ± 0.4 | — | 39.7 ± 0.4 |
| NGC 3516 | — | 17.2 ± 0.3 | — | 24.6 ± 0.3 | — | 40.5 ± 0.3 |
| NGC 3783 | >17.3 | 18.4 ± 0.3 | >24.2 | <25.5 | <40.5 | <41.9 |
| NGC 4051 | 15.3 ± 0.1 | 15.7 ± 0.2 | 23.6 ± 0.1 | 24.0 ± 0.2 | 40.7 ± 0.1 | 41.2 ± 0.2 |
| NGC 4151 | 15.1 ± 0.4 | <18.1 | >22.6 | 25.2 ± 0.6 | >40.7 | 43.4 ± 0.6 |
| NGC 4395 | — | 13.9 ± 0.4 | — | <21.5 | — | <36.1 |
| NGC 5506 | 14.0 ± 0.1 | 15.4 ± 0.5 | 23.9 ± 0.1 | 25.3 ± 0.5 | 43.3 ± 0.1 | 44.7 ± 0.5 |
| PDS 456 | 16.0 ± 0.1 | 18.2 ± 0.2 | 25.8 ± 0.1 | 27.9 ± 0.2 | 45.3 ± 0.1 | 47.5 ± 0.2 |
| SW J2127 | 13.9 ± 0.1 | 16.8 ± 0.5 | 23.4 ± 0.1 | 26.2 ± 0.5 | 42.8 ± 0.1 | 45.6 ± 0.5 |

2.4.2 Bolometric luminosity and Eddington ratio

We estimate the bolometric luminosity for each AGN: $L_{\text{bol}} = \kappa_{\text{bol}} L_{2-10\text{keV}}$, where $L_{2-10\text{keV}}$ is the unattenuated source luminosity integrated between 2 and 10 keV and κ_{bol} is the bolometric correction factor. Various studies have shown that the spread of κ_{bol} amongst individual AGN is quite large (e.g. Elvis et al. 1994), whilst the correction appropriate for a particular AGN can be a function of luminosity (Marconi & Hunt 2003; Hopkins, Richards & Hernquist 2007), Eddington ratio (Vasudevan & Fabian 2007; Lusso et al. 2010) or both (Lusso et al. 2010), such that the uniform application of a single bolometric correction factor may be inappropriate for a heterogeneously selected sample. We therefore primarily use the κ_{bol} values listed in the works of Vasudevan & Fabian (2007) and Vasudevan et al. (2010, 2009) which are empirically determined on the basis of the broad-band spectral energy distributions of each AGN. Several of the AGNs in the sample are not listed in these works. For these sources, we searched the literature for an appropriate bolometric luminosity and used that instead (see Table 1 caption). Only for 3C 111 and CBS 126 were we unable to locate either a bolometric correction factor or an definite and empirically measured bolometric luminosity; for these two sources, we simply assume that $\kappa_{\text{bol}} \sim 30$, which is similar to that found previously for other sources which harbour UFOs (e.g. Tombesi et al. 2012a). Values for $L_{2-10\text{keV}}$ (hereafter denoted L_{X}), κ_{bol} and the resultant estimate of L_{bol} are all listed in Table 3. For completeness, we also compute the likely Eddington ratio of each AGN: $\lambda \equiv L_{\text{bol}}/L_{\text{Edd}}$, where $L_{\text{Edd}} = 4\pi G m_{\text{p}} M_{\text{BH}} c \sigma_{\text{T}}^{-1} \simeq 1.26 \times 10^{38} (M_{\text{BH}}/M_{\odot})$.

3 DATA ANALYSIS

3.1 Preparing the data

Several of the AGNs in the sample have Fe K absorption detected in multiple epochs (see Paper I), such as is the case for 3C 111, Mrk 766, NGC 1365, NGC 3227, NGC 3783, NGC 4051, or have more than one absorption trough which comprises a multivelocity

system (PDS 456 and APM 08279+5255). We account for the intrinsic outflow variability between epochs for these AGN by folding the uncertainties associated with each individual measurement into the error bar. In the majority of cases, we find the individual measurements to be largely consistent within the errors, suggesting that the outflow is persistent across the different epochs. Only in 3C 111, Mrk 766 and NGC 4051 do there appear to be significant differences at the 90 per cent level (see Table 1), and this is reflected in the broad error bars for these sources. For PDS 456 and APM 08279+5255, which both appear to have two Fe K absorption systems at different velocities, we adopt the mean velocity and again fold the range of possible velocities into the error bar. Table 1 summarizes the measured parameters of the winds detected in the *Suzaku* sample.

3.2 Parameter constraints

Using the values in Table 1, we first computed the wind parameters in standard units and then normalized them to the appropriate values for a given BH mass; the resultant values are noted in Tables 2 and 3, respectively. We normalized the specific parameters as follows: (i) the distance (r_{max} , r_{min}) to units of Schwarzschild radius $r_{\text{s}} = 2GM_{\text{BH}}/c^2$, (ii) the kinetic luminosity ($L_{\text{w}}^{\text{max}}$, $L_{\text{w}}^{\text{min}}$) to L_{Edd} , (iii) the mass outflow rate ($\dot{M}_{\text{w}}^{\text{max}}$, $\dot{M}_{\text{w}}^{\text{min}}$) to the Eddington accretion rate $\dot{M}_{\text{Edd}} = L_{\text{Edd}}/\eta c^2$ (assuming $\eta = 0.06$ for the accretion efficiency) and (iv) the momentum rate to the Eddington momentum rate $\dot{p}_{\text{Edd}} = L_{\text{Edd}}/c$. Algebraically, this yields a set of eight normalized equations:

$$r_{\text{max}}/r_{\text{s}} = (2GM_{\text{BH}}\xi N_{\text{H}}n_{\text{e}})^{-1}n_{\text{H}}L_{\text{ion}}c^2, \quad (6a)$$

$$r_{\text{min}}/r_{\text{s}} = (c/v_{\text{w}})^2, \quad (6b)$$

$$\dot{M}_{\text{w}}^{\text{max}}/\dot{M}_{\text{Edd}} = (4\pi G\xi M_{\text{BH}}n_{\text{e}})^{-1}\Omega\sigma_{\text{T}}n_{\text{H}}L_{\text{ion}}v_{\text{w}}\eta c, \quad (6c)$$

$$\dot{M}_{\text{w}}^{\text{min}}/\dot{M}_{\text{Edd}} = (2\pi v_{\text{w}}n_{\text{e}})^{-1}\Omega\sigma_{\text{T}}n_{\text{H}}N_{\text{H}}\eta c, \quad (6d)$$

Table 3. Summary of normalized wind parameters.

| Source | $\log L_X^\dagger L_{\text{bol}}$ | κ_{bol} | $\log r_{\text{min}}$ (r_s) | $\log r_{\text{max}}$ (r_s) | $\log \dot{M}_w^{\text{min}}$ (\dot{M}_{Edd}) | $\log \dot{M}_w^{\text{max}}$ (\dot{M}_{Edd}) | $\log L_w^{\text{min}}$ (L_{Edd}) | $\log L_w^{\text{max}}$ (L_{Edd}) |
|--------------|-------------------------------------|-----------------------|-------------------------------------|-------------------------------------|---|---|---|---|
| 3C 111 | 44.35 ± 0.18 45.7 | 20 | 2.3 ± 0.1 | 3.7 ± 0.6 | −1.3 ± 0.1 | 0.1 ± 0.6 | −2.9 ± 0.1 | −1.5 ± 0.6 |
| 3C 390.3 | 44.42 ± 0.01 45.5 | 13.3 ³ | 1.7 ± 0.1 | <2.8 | >−1.1 | <−0.0 | >−2.1 | <−2.8 |
| 4C +74.26 | 44.95 ± 0.01 46.3 | — | 1.5 ± 0.1 | <5.0 | >−2.8 | 0.2 ± 0.5 | >−3.8 | −0.6 ± 0.5 |
| APM 08279 | 46.73 ± 0.03 47.4 | −1 | 1.1 ± 0.1 | 4.4 ± 0.3 | −1.9 ± 0.1 | 1.4 ± 0.3 | −2.3 ± 0.1 | 1.0 ± 0.3 |
| CBS 126 | 43.82 ± 0.02 45.1 | 20 | — | <2.6 | — | −1.1 ^{+0.2} _{−0.3} | — | −4.3 ^{+0.2} _{−0.3} |
| ESO 103-G035 | 43.47 ± 0.01 44.6 | 12.6 ⁴ | 2.5 ^{+0.5} _{−0.3} | <6.0 | >2.6 | 0.0 ± 1.1 | >−4.4 | −1.8 ± 1.1 |
| MCG -6-30-15 | 42.77 ± 0.01 44.0 | 16.1 ⁵ | 4.3 ^{+0.3} _{−0.2} | 5.23 ± 0.1 | −1.1 ^{+0.3} _{−0.2} | −0.2 ± 0.1 | −4.7 ^{+0.3} _{−0.2} | −3.8 ± 0.1 |
| MR 2251-178 | 44.61 ± 0.01 46.0 | 22.0 ³ | 1.7 ± 0.1 | 5.2 ± 0.1 | −2.53 ± 0.1 | 1.0 ^{+0.5} _{−0.8} | −3.55 ± 0.1 | 0.1 ^{+0.5} _{−0.8} |
| Mrk 279 | 42.78 ± 0.01 44.0 | 10.1 ⁵ | 1.3 ± 0.1 | 2.6 ^{+0.6} _{−0.4} | −1.36 ± 0.1 | −0.1 ^{+0.6} _{−0.4} | −1.98 ± 0.1 | −0.7 ^{+0.6} _{−0.4} |
| Mrk 766 | 42.73 ± 0.05 44.1 | 48.9 ⁵ | 2.8 ± 0.1 | 5.0 ± 0.3 | −1.3 ± 0.1 | 0.9 ± 0.3 | −3.4 ± 0.1 | −1.2 ± 0.3 |
| NGC 1365 | 42.32 ± 0.02 44.1 | 101.6 ⁵ | — | 2.1 ± 0.1 | — | −1.7 ± 0.1 | — | −5.2 ± 0.1 |
| NGC 3227 | 42.13 ± 0.06 43.3 | 15.3 ⁵ | — | 2.9 ± 0.3 | −0.4 ^{+1.4} _{−0.5} | −2.2 ± 0.3 | — | −6.1 ± 0.3 |
| NGC 3516 | 43.07 ± 0.01 44.1 | 14.0 ⁵ | — | 4.1 ± 0.3 | — | −1.13 ± 0.3 | — | −5.23 ± 0.3 |
| NGC 3783 | 43.14 ± 0.05 44.5 | 23.7 ⁵ | >4.3 | 5.4 ± 0.3 | >−1.5 | <−0.1 | <−5.1 | <−4.3 |
| NGC 4051 | 41.65 ± 0.04 43.2 | 41.1 ⁴ | 3.5 ± 0.1 | 3.9 ± 0.2 | −0.9 ± 0.1 | −0.4 ± 0.2 | −3.7 ± 0.1 | −3.2 ± 0.2 |
| NGC 4151 | 42.34 ± 0.02 43.6 | 18.3 ⁴ | 2.5 ^{+0.5} _{−0.3} | <5.6 | >−2.7 | 0.0 ± 0.6 | >−4.5 | −1.8 ± 0.6 |
| NGC 4395 | 40.38 ± 0.01 41.7 | 22.5 ³ | — | 3.8 ± 0.4 | — | <−1.4 | — | <−7.5 |
| NGC 5506 | 43.19 ± 0.01 44.4 | 16.8 ⁵ | 1.2 ± 0.1 | 2.6 ± 0.5 | −1.6 ± 0.1 | −0.2 ± 0.5 | −2.1 ± 0.1 | −0.7 ± 0.5 |
| PDS 456 | 44.62 ± 0.01 47.0 | −2 | 1.1 ± 0.1 | 3.3 ± 0.2 | −1.8 ± 0.1 | 0.4 ± 0.2 | −2.2 ± 0.1 | 0.0 ± 0.2 |
| SW J2127 | 43.15 ± 0.01 44.5 | — | 1.3 ± 0.1 | 4.1 ± 0.5 | −2.0 ± 0.1 | 0.8 ± 0.5 | −2.6 ± 0.1 | 0.3 ± 0.5 |

Notes. ¹The bolometric luminosity of APM 08279+5255 is taken as $L_{\text{bol}} = 7 \times 10^{15} \mu_L^{-1} L_\odot$ (Lewis et al. 1998; Riechers et al. 2009; Saez & Chartas 2011), assuming a conservative magnification factor of $\mu_L = 100$ (Egami et al. 2000); ² L_{bol} for PDS 456 taken from Reeves & Turner (2000); ^{3,4,5} κ_{bol} correction factors taken from Vasudevan & Fabian (2007, 2009) or Vasudevan et al. (2010), respectively. [†] L_X here corresponds to the unattenuated source luminosity between 2 and 10 keV.

$$L_w^{\text{max}} / L_{\text{Edd}} = (4\pi G M_{\text{BH}} c \xi n_e)^{-1} \Omega \sigma_{\text{T}} n_{\text{H}} L_{\text{ion}} v_w^3, \quad (6e)$$

$$L_w^{\text{min}} / L_{\text{Edd}} = (4\pi c n_e)^{-1} \Omega \sigma_{\text{T}} n_{\text{H}} N_{\text{H}} v_w, \quad (6f)$$

$$\dot{p}_w^{\text{max}} / \dot{p}_{\text{Edd}} = (4\pi G M_{\text{BH}} c \xi n_e)^{-1} \Omega \sigma_{\text{T}} n_{\text{H}} L_{\text{ion}} v_w^2, \quad (6g)$$

$$\dot{p}_w^{\text{min}} / \dot{p}_{\text{Edd}} = (4\pi n_e)^{-1} \Omega \sigma_{\text{T}} N_{\text{H}}. \quad (6h)$$

Errors on the normalized parameters were determined by both propagating the uncertainties on the measured XSTAR values through the various equations, and also taking into account the error on M_{BH} where applicable. For CBS 126, NGC 1365, NGC 3227, NGC 3516 and NGC 4395, which all have relatively slow wind velocities of the order of $v_w < 0.01c$ (see Table 1), we found that $r_{\text{min}} > r_{\text{max}}$ (within the errors) which indicates that either: (i) the outflow may not have achieved the requisite escape speed for its observed location, in which case the wind material may not escape into the host galaxy unless it is subject to an additional acceleration mechanism once it has crossed the LOS, or (ii) that the wind stream line flows almost perpendicularly along the LOS such that we are only seeing the tangential component of v_w . While the associated error bars in these slow sources may overlap (suggesting that the upper and lower limits may be formally consistent) we choose to only report the upper limiting quantities for these five AGNs in Tables 2 and 3 so that we do not introduce an artificially tight (and potentially misleading) constraint on the properties of the wind. In the remainder of our analysis, we separate the outflow sample into two groups based on their velocity: those with $v_w \leq 0.01c$ (3000 km s^{−1})

are hereafter classified as ‘slow’, whilst those with $v_w > 0.01c$ are ‘fast’. Whilst the chosen velocity threshold is essentially arbitrary, this classification scheme provides a useful means to distinguish between those pristine ‘disc-winds’ from those which may just be the higher ionization component of the inner-BLR or the more distant warm absorber, with which they probably share an overlap in velocity space and location (e.g. see Tombesi et al. 2013).

In Fig. 1, we show constraint diagrams for the various wind parameters.² From Figs 1(a) and (b), it is clear that the absorbing material is distributed across several orders of magnitude in distance from their BH, with r spanning $\sim 10^{14-20}$ cm or $\sim 10^{1-6} r_s$. By considering the mean range,³ we see that the fast outflows appear to be located between $\langle r \rangle \sim 10^{15-18}$ cm ($\sim 10^{2-4} r_s$) distance range. This corresponds to ~ 0.0003 – 0.3 pc and implies that whilst there may be some limited overlap between the fast Fe K outflows and the traditional soft X-ray warm absorber (which is usually inferred to be on parsec-scale distances, e.g. Crenshaw, Kraemer & George 2003; Blustin et al. 2005; Kaastra et al. 2012 and references therein), the former are generally located much closer to their central BH. This is in agreement with other results in the literature which find that the highly ionized Fe K outflows most likely originate in a wind from the inner regions of the accretion disc (e.g. Proga & Kallman 2004; Schurch, Done & Proga 2009; Sim et al. 2008; King 2010),

² Note that we do not show a constraint diagram for \dot{p}_w because it is simply proportional to L_w/v_w and therefore follows a similar overall distribution to the one for L_w .

³ Here, ‘mean range’ refers to the range spanning between the means for the upper and lower limit values, respectively, i.e. it corresponds to the range between $\langle r_{\text{max}} \rangle$ and $\langle r_{\text{min}} \rangle$ for the fast outflows (see the caption for more details).

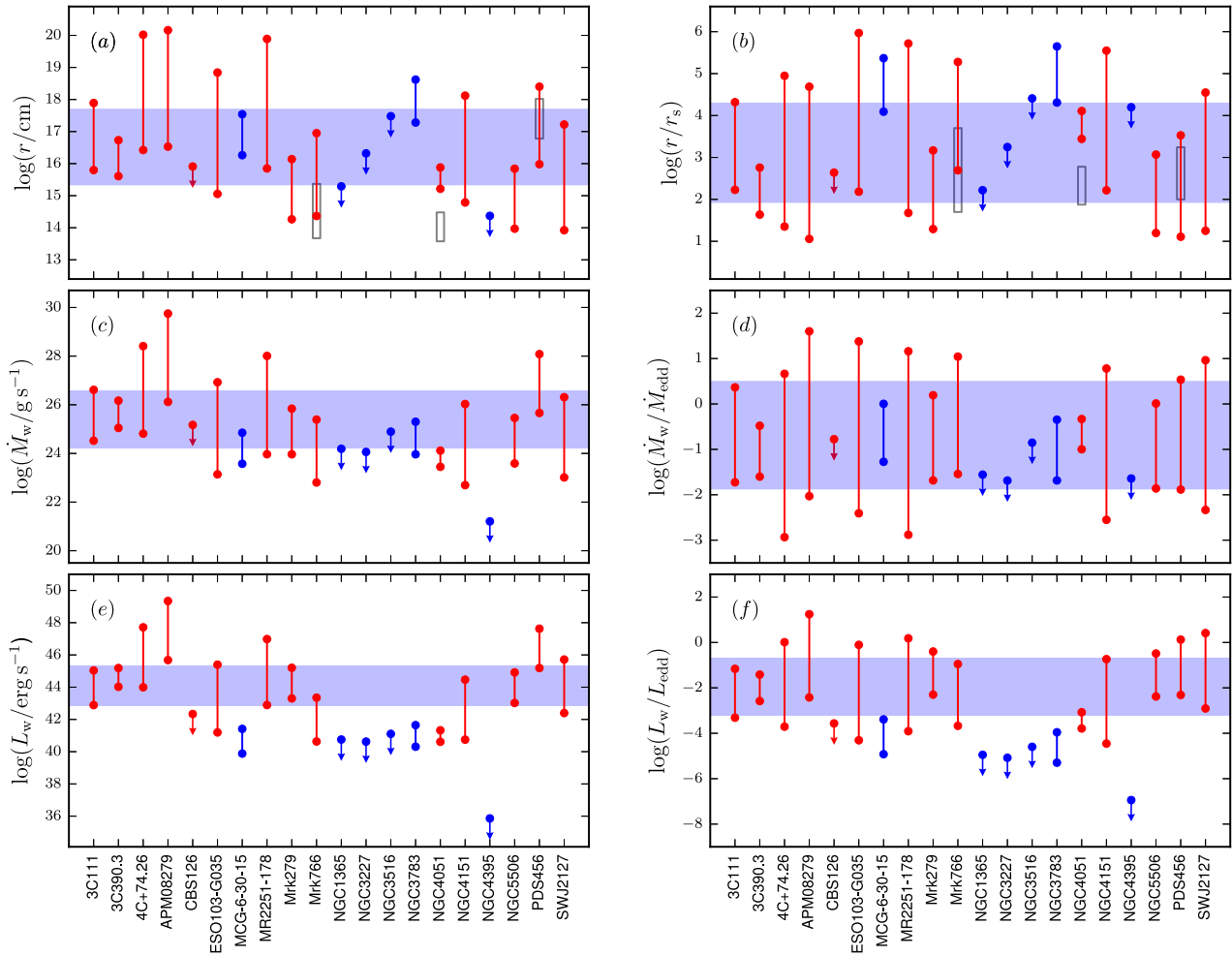


Figure 1. Constraint diagram showing the parameter ranges occupied by the *Suzaku*-detected Fe K outflows. The left-hand panels show the raw limiting values (taken from Table 2) and the right-hand panels show the normalized ones (from Table 3). The top, middle and bottom rows are the constraints on distance, mass outflow rate and kinetic power. In all panels, the red and blue lines show sources with $v_w > 3000 \text{ km s}^{-1}$ and $v_w < 3000 \text{ km s}^{-1}$, respectively. For Mrk 766, NGC 4051 and PDS 456, the grey boxes show the range in location measured by Risaliti et al. (2011), Miller et al. (2007, 2010), Turner et al. (2007) and Gofford et al. (2014) on the basis of variability or X-ray reverberation studies, respectively (see text for further details). The blue/shaded areas show the ‘mean range’, spanning the mean of all upper and all lower limiting values of the fast winds only.

and is also consistent with what Tombesi et al. (2012a) found on the basis of the *XMM-Newton* outflow sample.

Following our estimates on r , Fig. 1(c) shows an analogous plot for \dot{M}_w . The constraints on \dot{M}_w also vary significantly, spanning almost eight orders of magnitude, and ranging from $<10^{22} \text{ g s}^{-1}$ in the dwarf Seyfert galaxy NGC 4395 all the way up to $\sim 10^{30} \text{ g s}^{-1}$ in the massive quasar APM 08279+5255. The other winds have \dot{M}_w estimates distributed between these two extremes, with the mean range for the fast outflows falling between $\langle \dot{M}_w \rangle \sim 10^{24-26} \text{ g s}^{-1}$ ($\sim 0.01\text{--}1 M_\odot \text{ yr}$). Interestingly, and while there remains considerable uncertainty in each case, $\langle \dot{M}_w^{\text{max}} \rangle$ is of the order of the Eddington rate, i.e. $\langle \dot{M}_w^{\text{max}} \rangle \approx \dot{M}_{\text{Edd}}$, while even $\langle \dot{M}_w^{\text{min}} \rangle$ is still ~ 1 per cent of \dot{M}_{Edd} . This immediately suggests that the winds are transporting a substantial amount of material into their host galaxies. The ejected mass has a large kinetic power: Fig. 1(e) shows that while the range of L_w again spans several orders of magnitude, the overall mean for the fast flows is constrained to between $\langle L_w \rangle \sim 10^{43}$ and $10^{45} \text{ erg s}^{-1}$. The normalized values show that, on average, $\langle L_w \rangle$ is $< L_{\text{Edd}}$, with a mean range of $\langle L_w \rangle \sim (0.1\text{--}10 \text{ per cent}) L_{\text{Edd}}$. The wind mass outflow rate and the associated kinetic power are therefore significant

fractions of the Eddington-limited values. In Section 4.3, we extend these results to assess whether the wind is energetically significant in terms of feedback.

As mentioned previously, our estimates on r are subject to large uncertainties. It can therefore be useful to compare a few of our inferred values with some more robust measurements available in the literature. We consider here the exemplar cases of PDS 456 and Mrk 766. Currently, PDS 456 is the only AGN where the location of the Fe K-shell wind has been robustly constrained from discrete line variability. On the basis of a recent *Suzaku* campaign, Gofford et al. (2014) constrained the Fe K absorber in PDS 456 to $\sim 100\text{--}1700 r_s$. This is entirely consistent with what we infer using equations (1) and (2). Similar is also true for Mrk 766: Miller et al. (2007) and Turner et al. (2007) estimate the absorbing gas to be located on distances of the order of $30\text{--}50 r_s$, while Risaliti et al. (2011) later showed that both the broad-band spectral variability and the Fe K absorption could also be attributed to a stratified, inhomogeneous and clumpy absorber which occults the AGN at a distance of $\sim 500\text{--}5000 r_s$. These estimates are again consistent with what we find here (see Figs 1a and b), suggesting that the

two studies are perhaps probing the same layer of gas. Elsewhere in the literature, Miller et al. (2010) used a detailed spectral timing analysis to deconvolve the X-ray spectrum of NGC 4051 into its constituent absorption- and reflection-dominated components, with the hard X-ray reverberation signal suggesting that the reflecting gas is $\sim 100\text{--}600 r_g$ from the BH (at the 90 per cent level, see Miller et al. 2010). This is around one order of magnitude closer than where we estimate the Fe K absorbing gas to be in this source. One possibility is that the reflecting layer detected by Miller et al. (2010) is physically distinct to the Fe K absorber, possibility associated with the Compton-thick occulting clouds posited by Tatum et al. (2013, in preparation).

3.3 Correlations with bolometric luminosity

Supplementing the constraint diagrams, we also carried out a correlation analysis. This analysis bears notable similarities to the one recently conducted by Tombesi et al. (2013) for the *XMM-Newton* outflow sample. There, the authors combined the parameters of the hard-band absorber, as measured in their Fe K absorption survey with *XMM-Newton*, with those for the soft X-ray warm absorber collated from the literature to determine how the AGN wind parameters vary globally in relation to their distance from the BH (Tombesi et al. 2013). Here, noting that *Suzaku* does not have the soft X-ray energy resolution necessary to constrain the velocity of the soft X-ray absorbing gas, we consider only the Fe K absorbers and do not consider the warm absorber. Whilst this leads to a relatively small dynamic range in terms of their distance from the BH, the heterogeneous nature of the *Suzaku* sample encompasses a significantly broader range of AGN in terms of their bolometric luminosity (L_{bol}). The *Suzaku* sample therefore offers a useful opportunity to probe how the wind parameters may vary in comparison to the luminosity of their host AGN.

We compute linear regressions of the form $\log(y) = \alpha + \beta \log(x)$, where α and β are the intercept and slope of the straight line fit, respectively, using the `linmix_err` Bayesian regression algorithm⁴ (Kelly 2007). This routine employs Markov Chain Monte Carlo (MCMC) techniques to self-consistently account for measurement errors and intrinsic scatter in the data, whilst also allowing for limited censorship (upper limits only) in the independent variable. We computed regression parameters from 100 000 MCMC realizations of the data by the `linmix_err` routine. This yields posterior distributions for the intercept α , slope β , standard deviation (scatter) of the data σ and the Pearson linear correlation coefficient R_p . We adopt the median value of the posterior distribution as our ‘best fit’ to the data, and estimate the parameter errors as the range of simulated values which encompass 90 per cent of the posterior distribution about the median.

We note the slow winds have $v_w \leq 3000 \text{ km s}^{-1}$ which is similar to that measured in the traditional soft X-ray warm absorber. It is therefore unclear whether these winds are associated with a bona fide disc-wind or alternatively with another layer of gas which is more distant from the BH, e.g. with a higher ionization component of the warm absorber. For this reason, we fit all subsequent regression analyses to the fast winds only, $v_w > 3000 \text{ km s}^{-1}$. Owing to their high velocities, these faster winds are kinematically distinct from the soft X-ray absorbing gas and thus may more obviously represent the signature of a pristine disc-winds. This limits our analysis

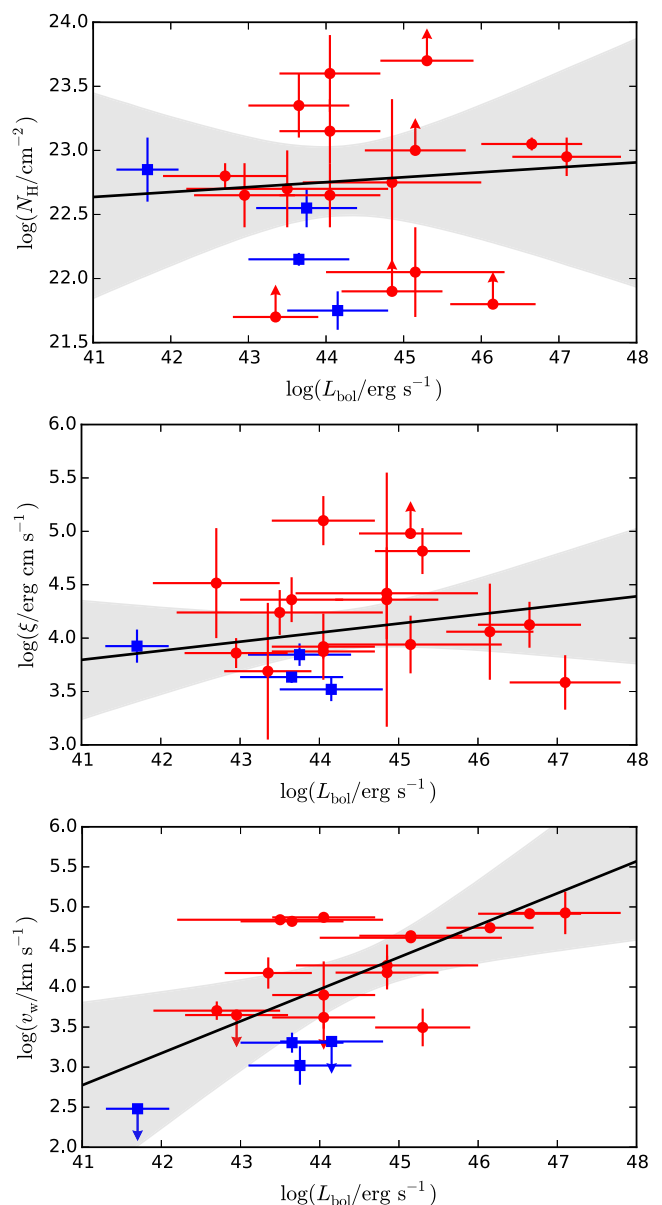


Figure 2. Scatter plots showing $\log L_{\text{bol}}$ versus: (a) $\log(N_{\text{H}}/\text{cm}^{-2})$, (b) $\log(\xi/\text{erg cm s}^{-1})$ and (c) $\log(v_w/\text{km s}^{-1})$, as noted in Table 2. The red and blue data points correspond to AGN with $v_w > 3000 \text{ km s}^{-1}$ and $v_w < 3000 \text{ km s}^{-1}$, respectively. The solid black line corresponds to the ‘best-fitting’ linear regression to the fast winds, as estimated from 100 000 MCMC realizations of the data set using the `linmix_err` Bayesian regression algorithm (Kelly 2007), while the grey shaded area denotes the region which contains 90 per cent of the posterior probability distribution. The parameters of the fits are reported in Table 4.

to the winds which are capable of escaping the gravitational potential of the central BH (based on our LOS), and are therefore more likely to have an effect on the host galaxy in terms of feedback.

In Fig. 2, we show how $\log(N_{\text{H}}/\text{cm}^{-2})$, $\log(\xi/\text{erg cm s}^{-1})$ and $\log(v_w/\text{km s}^{-1})$, as measured in Paper I and summarized in Table 2, vary with L_{bol} . From Figs 2(a) and (b), we find no discernible relationships between either N_{H} or ξ and L_{bol} ($\beta = 0$ at 90 per cent confidence): the $\log(N_{\text{H}}/\text{cm}^{-2})$ values cover a wide range for a given value of L_{bol} , while $\log(\xi/\text{erg cm s}^{-1})$ appears to cluster at around ~ 4 , mirroring the median of the parameter distributions that we found in Paper I. This also appears to be true for both the fast

⁴ Available in IDL from

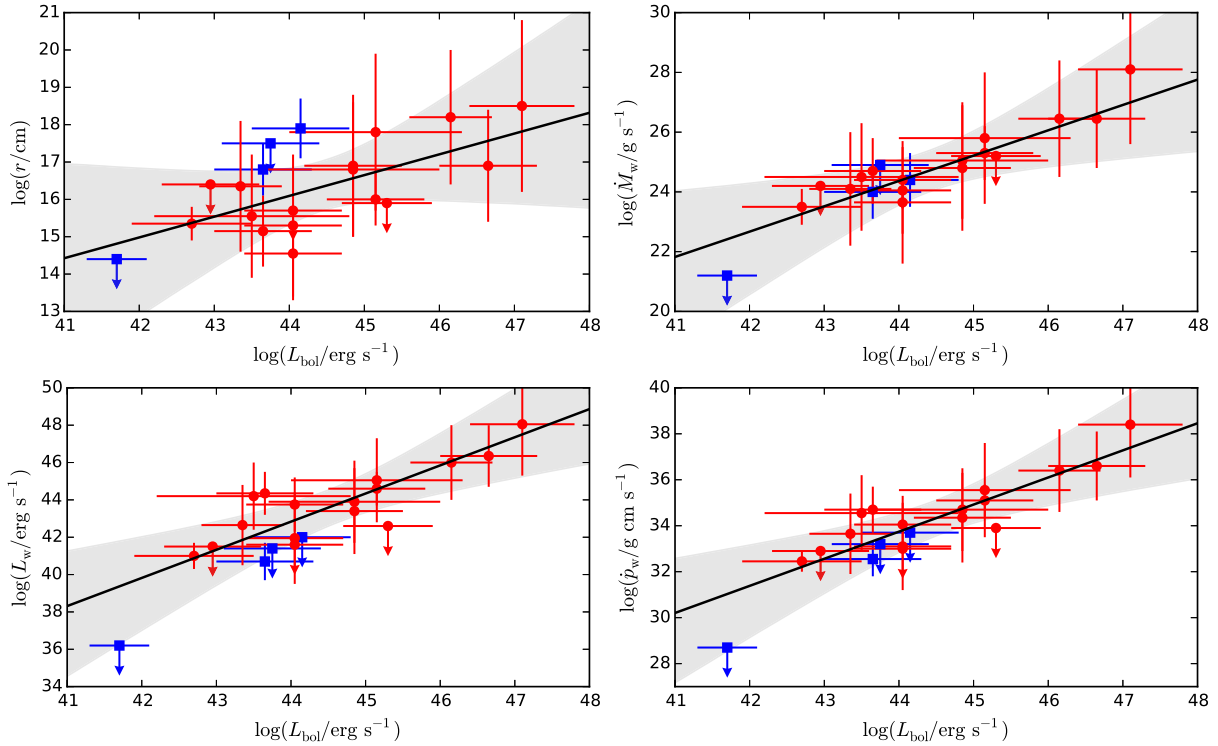


Figure 3. Scatter plots showing $\log L_{\text{bol}}$ against the raw wind parameters: (a) $\log(r/\text{cm})$, (b) $\log(\dot{M}_w/\text{g s}^{-1})$, (c) $\log(L_w/\text{erg s}^{-1})$ and (d) $\log(\dot{p}_w/\text{g cm s}^{-1})$. Data points and plot components have the same meaning as in previous figures.

and the slow wind sub-samples. Conversely, there is a correlation present between v_w and L_{bol} , with the faster winds tending to be observed in more luminous AGN. The fast winds have a constrained slope of $\beta = 0.4^{+0.3}_{-0.2}$, with $v_w \propto L_{\text{bol}}^\beta$ and $\beta \approx 0.5$. In contrast, the slower systems, i.e. those with $v_w < 0.01c$, appear to be isolated to the lower-left quarter of the plot which suggests that low-velocity winds are preferentially located in lower luminosity systems. This is consistent with the overall picture that wind velocity is correlated with L_{bol} ; fitting a correlation to the entire sample yields $\beta_{\text{all}} = 0.5^{+0.4}_{-0.2}$. This is formally consistent within the errors but does suggest that the inclusion of the slow winds may skew the relationship to a slightly steeper slope.

Fig. 3 then shows how the raw wind parameters vary with L_{bol} . Here, and in subsequent figures, each data point corresponds to the mid-point in the range for each parameter, calculated as e.g.

$$r_{\text{mid}} = \frac{(r_{\text{max}} + \text{perr}) + (r_{\text{min}} - \text{nerr})}{2}, \quad (7)$$

where **nerr** and **perr** are the negative and positive error, respectively. The associated error bars then denote the range between the maximum and minimum values, including the uncertainties arising from the input variables. In equation (7), we are essentially only considering the upper and lower limits for *all* derived values in the sample, such that in the cases where the minimum and maximum values are themselves unconstrained (i.e. only an upper or lower limit is available with no associated error) we are able to use the censored value when calculating the mid-point. This ensures that all of the available data are included in our analysis.

To first order, we find that all of the raw wind parameters are plausibly correlated with L_{bol} to some degree. Whilst the slope for $L_{\text{bol}}-r$ is tentative, and marginally consistent with $\beta = 0$ at the 90 per cent level (Fig. 3, the overall distribution is skewed towards a

positive relationship among the data, which suggests that the winds are observed at larger radii in brighter AGN, as would be expected given the larger typical size-scales in these systems). Fig. 3 then shows that there is a positive relationship between L_{bol} and \dot{M}_w ($\beta = 0.9^{+0.8}_{-0.6}$), such that the winds in more luminous AGN are correspondingly more massive. In both of these cases, the slow winds appear to follow the same overall trend as the fast ones, despite not formally being included in the regression computation. Similar is also true for $L_{\text{bol}}-L_w$ and $L_{\text{bol}}-\dot{p}_w$, which share strong positive slopes of $\beta = 1.5^{+1.0}_{-0.8}$ ($R_p \geq 0.54$) and $\beta = 1.2^{+0.8}_{-0.7}$ ($R_p \geq 0.55$), respectively. The observed slopes here likely stem from their mutual dependence on v_w : $L_w \propto v_w^3$ and $\dot{p}_w \propto v_w^2$. Indeed, the $L_w \propto L_{\text{bol}}^{1.5}$ and $\dot{p}_w \propto L_{\text{bol}}^{1.0}$ relationships implied here are formally consistent with the $v_w \propto L_{\text{bol}}^{0.5}$ found earlier in Fig. 2(c). This ultimately suggests that the wind velocity may be the driving factor in the observed relationships, with the uncertainties on the other wind parameters only adding to the underlying scatter of the data. In conclusion, these results indicate that disc-winds in higher luminosity sources are (tentatively) located farther away from their central BH in terms of absolute distance (as expected for larger M_{BH}), they contain more material, and they are accordingly more energetic. The driving factor behind these relationships appears to be the wind velocity, which is faster in more luminous objects.

Similar scatter plots for the normalized wind parameters are shown in Fig. 4. In contrast to before, this time we are unable to discern any linear correlations in the normalized $L_{\text{bol}}-(r/r_s)$ and $L_{\text{bol}}-(\dot{M}_w/\dot{M}_{\text{Edd}})$ plots. In fact, and whilst there is a weak negative correlation suggested in the case of the former, both of the relationships are formally consistent with $\beta = R_p = 0$ such that the r/r_s and $\dot{M}_w/\dot{M}_{\text{Edd}}$ ratios are largely independent of L_{bol} . The fast winds are therefore observed at similar relative distances from the SMBH regardless of the source luminosity, and their mass outflow rate is a

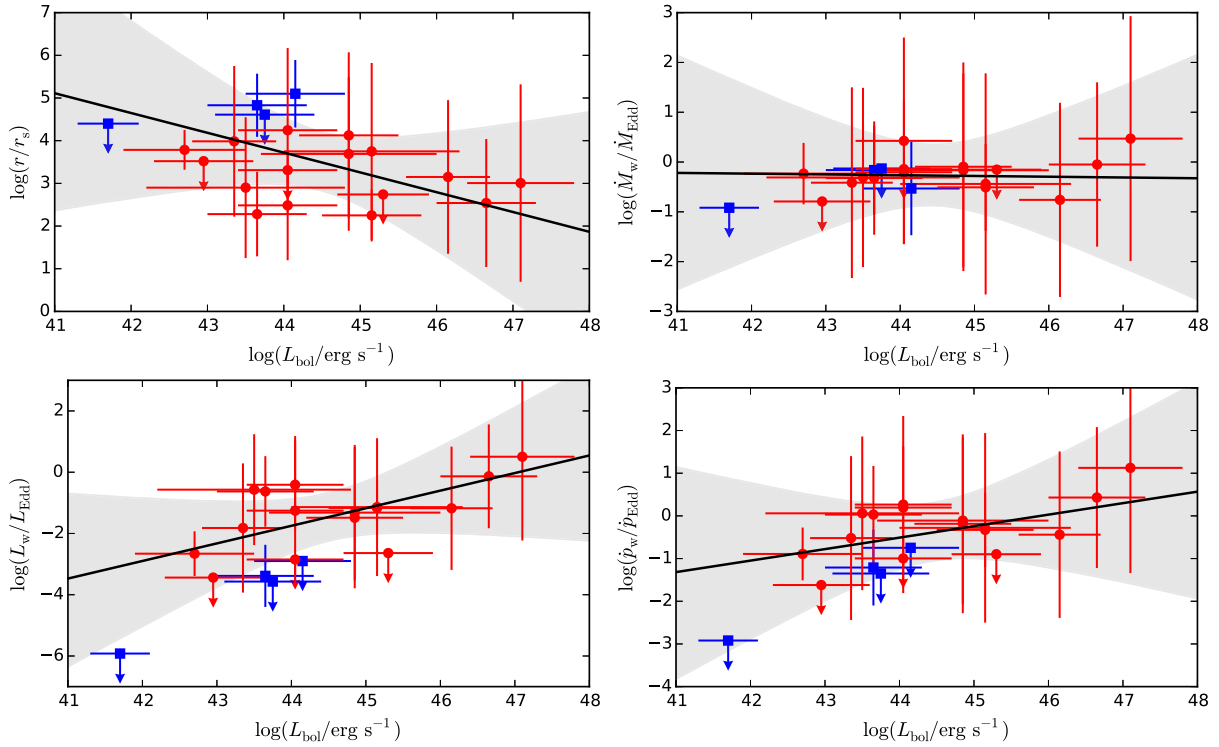


Figure 4. More scatter plots, this time showing $\log L_{\text{bol}}$ against the normalized wind parameters: (a) $\log(r/r_s)$, (b) $\log(\dot{M}_w/\dot{M}_{\text{Edd}})$, (c) $\log(L_w/L_{\text{Edd}})$ and (d) $\log(\dot{p}_w/\dot{p}_{\text{Edd}})$. Data points and plot components have the same meaning as in previous figures.

similar fraction of \dot{M}_{Edd} . The slow winds also appear to follow consistent relationships. Both L_w/L_{Edd} (Fig. 4d) and $\dot{p}_w/\dot{p}_{\text{Edd}}$ (Fig. 4e) are also formally consistent with a slope of zero, although weak positive correlations are suggested. In both Figs 4(c) and (d), the slow winds appear to lie below the computed regression line for the fast winds. Their location in these plots appears to mirror that seen in Fig. 4 for the $L_{\text{bol}}-v_w$ correlation, which is again consistent with the wind velocity being the parameter behind the relationships.

For completeness, we also tried to determine whether the parameters of the wind scale with the Eddington ratio of the AGN. Unfortunately, however, the relatively tight clustering in the data at $\lambda = 0.1$, compounded by the lack of Fe K wind parameters in low Eddington ratio sources (i.e. $\lambda \lesssim 10^{-2}$) sources, meant that we were unable to discern any plausible relationships among the data. As a consistency check, we also checked for evidence of relationships amongst the raw L_X values finding that the same correlations are still present in the data, albeit at lower significance level. The presence of the same correlations in both the contiguous L_X and the non-simultaneous L_{bol} indicates that the correlations are not driven by uncertainties in κ_{bol} .

4 DISCUSSION

4.1 Comparisons with previous work

4.1.1 The XMM–Newton outflow sample

In Paper I, we showed that ~ 40 per cent of the AGN in the *Suzaku* sample exhibit evidence for highly ionized winds in their Fe K band. This is in agreement with the detection fraction reported by Tombesi et al. (2010a) for *XMM–Newton*, and corroborates their

conclusion that such winds are either (i) persistent over the active phase and thus have a large covering fraction, or (ii) transient in nature and have a covering fraction of unity for only a fraction of the time. We also showed that the $\log(N_{\text{H}}/\text{cm}^{-2})$ and $\log(\xi/\text{erg cm s}^{-1})$ distributions are also entirely consistent between the two samples, covering ranges of $21 \lesssim \log(N_{\text{H}}/\text{cm}^{-2}) \leq 24$ and $2 \lesssim \log(\xi/\text{erg cm s}^{-1}) \leq 6$ in both, respectively, with the outflow velocity v_w also being largely consistent on the basis of a Kolmogorov–Smirnov (KS) test. By comparing the distributions of wind parameters found in this work to those determined by Tombesi et al. (2010a), we similarly find that the Fe K absorbers detected in both the *Suzaku* and *XMM–Newton* studies are typified by the same range of physical parameters and occupy the same parameter space in terms of their overall location and energetics: $1 \lesssim \log(r/r_s) \lesssim 3$ (~ 0.001 –1 pc), $-2 \lesssim \log(\dot{M}_w/\dot{M}_{\text{Edd}}) \lesssim 0$ (~ 0.001 –10 M_{\odot} yr) and $-3 \lesssim \log(L_w/L_{\text{Edd}}) \lesssim 0$. Their medians are also similar with $\log(r/r_s) \sim 16(17)$, $\log(\dot{M}_w/\text{g s}^{-1}) \sim 25(25)$ and $\log(L_w/\text{erg s}^{-1}) \sim 44(44)$ for *Suzaku* (*XMM–Newton*), respectively, in natural units.

The *Suzaku* outflow sample therefore robustly confirms that Fe K absorption are a real physical component of emergent X-ray spectrum, that the implied winds likely have a large covering fraction and that they are typified by large column densities and high ionization. The range of outflow velocities, which are typically $v_w \geq 10\,000 \text{ km s}^{-1}$, but can be as low as $\lesssim 1000 \text{ km s}^{-1}$ in some cases, also suggests that the highly ionized winds share an overlap in velocity space with the traditional soft X-ray warm absorber. This is in line with the results of Tombesi et al. (2013) who argue that some of the soft X-ray warm absorbers could be the artefacts of accretion disc winds which have propagated farther away from the SMBH.

4.1.2 Comparison with King et al. (2013)

Another interesting study that we can consider is that of King et al. (2013), who compare the kinetic power of both the warm absorber and relativistic jets to L_{bol} of their respective source. They consider both black hole binaries (BHBs) and AGN which means that their study samples an extremely broad spectrum of both BH mass [$0.8 \lesssim \log(M_{\text{BH}}/M_{\odot}) < 10$] and bolometric luminosity [$37 \lesssim \log L_{\text{bol}} < 47$]. Even though the King et al. (2013) study deals primarily with the low-velocity soft X-ray warm absorber they do include a few ‘ultrafast outflows’ (which they define as those with $v_w \geq 0.01c$, as we have done in this work) as a matter of comparison, although they do not fit them directly as part of their regression analysis; two of their ‘ultrafast’ sources, namely 3C 111 and APM 08279+5255, are also included in our *Suzaku* sample. In principle, we could compare the properties of the *Suzaku*-detected outflows with those of King et al. (2013) to determine if our high velocity Fe K winds scale in a similar manner to the warm absorber. However, this is not possible in practice because the two studies compute the wind kinetic power according to different assumptions (King et al. (2013) give their quantities in units of per-covering-fraction), which means the intercept of any linear regressions will therefore be intrinsically offset. Even so, the *slope* of any linear fit – which is the important parameter as it measures how the kinetic power varies with bolometric luminosity – can still be compared because it is largely independent of the normalization.

King et al. (2013) find the kinetic power of warm absorber to be strongly correlated with L_{bol} , scaling with a tight global slope of $\log L_w \propto (1.58 \pm 0.07) \log L_{\text{bol}}$, while they found that those in AGN have a much flatter local slope of 0.63 ± 0.30 when considered in isolation (although that this could be driven by there being a smaller number of AGNs). In Fig. 3(d), we showed that kinetic power of the *Suzaku*-detected Fe K winds are also strongly correlated with the bolometric luminosity with a slope in the range $1.5^{+1.0}_{-0.8}$ ($R_p \geq 0.54$). This result, which is driven by the $v_w \propto L_{\text{bol}}^{0.5}$ dependence, is in good agreement with the slope reported by King et al. (2013) for their global sample and suggests that the fast Fe K winds may scale with L_{bol} in a similar manner to the warm absorber. To summarize, by comparing our results for those obtained by King et al. (2013), we see that the fast Fe K winds ($v_w > 0.01c$) is similar to that found for slow warm absorbers.

4.2 Are the winds radiatively accelerated?

In this work, we shown that the observed wind velocity, and subsequently the overall wind energetics, are proportional to the AGN bolometric luminosity. It is therefore tempting to conclude that the winds are accelerated by radiation pressure. In this section, we investigate whether radiation pressure alone can account for the observed correlations. In contrast to the undoubtedly line-driven winds which are prevalent in the UV spectra of AGNs, the high-ionization state of the Fe K winds suggests that line-driving is unlikely to be the dominant acceleration mechanism here. An alternative means of radiatively accelerating high-ionization gas is through Thomson/Compton scattering of the continuum X-ray photons. While line-driving may play a role in large BHs accreting near the Eddington limit (Hagino et al. 2015), we concentrate on the latter mechanism here.

In comparison to fast line-driven winds in the UV, the comparatively small interaction cross-section in highly ionized material means that high-ionization winds, such as those considered here, couple much less efficiently with the incident radiation and, as a

result, require either much larger source luminosities or higher column densities to achieve equivalent outflow velocities. A general characteristic of these ‘continuum-driven winds’ is that they are accelerated by photons scattering off free electrons in the absorbing gas which transfers a portion of the photon momentum to the material, hence causing a wind. This scenario has been considered extensively in the literature (e.g. King & Pounds 2003; King 2005, 2010; Reynolds 2012; Costa, Sijacki & Haehnelt 2015). We outline the general theory again here in an effort to search for relevant relationships with L_{bol} which we can relate to the results of our regression analysis.

The efficiency through which the photon momentum of the incident radiation field is transferred into the wind depends upon the plasmas optical depth to electron scattering, τ , which determines how many times each electron is likely to scatter in the flow. The total electron scattering optical depth of a homogeneous wind viewed from infinity down to radius r is given by

$$\tau = N_H \sigma_T = \sigma_T \int_r^\infty n dr = \frac{\sigma_T \dot{M}_w}{4\pi b m_p v_w r}, \quad (8)$$

where σ_T is the Thomson cross-section for electron-scattering, $n = \dot{M}_w / 4\pi b m_p v_w r^2$ is the electron number density in a mass-conserving spherical flow (cf. Section 2.2) and b is a geometrical factor which takes into account the global covering of the gas. Normalizing to the Eddington mass-accretion rate, $\dot{M}_{\text{Edd}} = L_{\text{Edd}} / \eta c^2 = 4\pi G M_{\text{BH}} m_p / \sigma_T \eta c$, then yields

$$\tau = \frac{1}{b\eta c} \frac{1}{v_w} \frac{G M_{\text{BH}}}{R} \frac{\dot{M}_w}{\dot{M}_{\text{Edd}}}, \quad (9)$$

as has been shown by King & Pounds (2003, see also King 2010 for a recent review). If we further assume that the measured flow velocity roughly corresponds to the escape velocity at the radius where the wind was launched, i.e. $R = R_{\text{esc}} = 2GM_{\text{BH}}/v^2$, we have

$$\tau = \frac{1}{2b\eta} \frac{v_w}{c} \frac{\dot{M}_w}{\dot{M}_{\text{Edd}}}, \quad (10)$$

which, after substituting for $\dot{M}_{\text{Edd}} = L_{\text{Edd}} / \eta c^2$ and re-arranging, becomes

$$\dot{p}_w = 2b\tau \frac{L_{\text{Edd}}}{c} = 2b \frac{\tau}{\lambda} \frac{L_{\text{bol}}}{c}, \quad (11)$$

where $\dot{p}_w \equiv \dot{M}_w v_w$, the geometric factor $b = \Omega / 4\pi$ and we have substituted from $\lambda = L_{\text{bol}} / L_{\text{Edd}}$ in the final step to express the relation in terms of observable quantities. Equation (11) therefore shows that a wind accelerated by electron scattering should have an outward momentum-rate which is proportional to that of the incident radiation field. This result (and the associated derivation) is directly equivalent to the one presented by King & Pounds (2003). Note that while τ has thus far been described as the optical depth to electron scattering it can in principal also account for additional sources of opacity in the flow, e.g. though bound–free or bound–bound absorption, which can also be further boosted by multiple scattering events. A reasonable way of testing for the continuum-driven wind scenario is to therefore see whether \dot{p}_w and $\dot{p}_{\text{bol}} \equiv L_{\text{bol}}/c$ follow a linear trend. Fig. 5(a) shows that the two are clearly correlated, with a slope of $\beta = 1.2^{+0.8}_{-0.7}$. This is consistent with the order of unity expected from equation (11), and is also agreement with the slope of $\beta_{\text{xmm}} \sim 1.6 \pm 1.1$ found by Tombesi et al. (2013). Fig. 5(b) is a constraint diagram for $\log(\dot{p}_w / \dot{p}_{\text{bol}})$ showing the deviations from the expected ratio of unity; in total, most of the sources (13/20; 65 per cent of the sample) are formally consistent with the ratio of unity expected for a continuum-driving scenario.

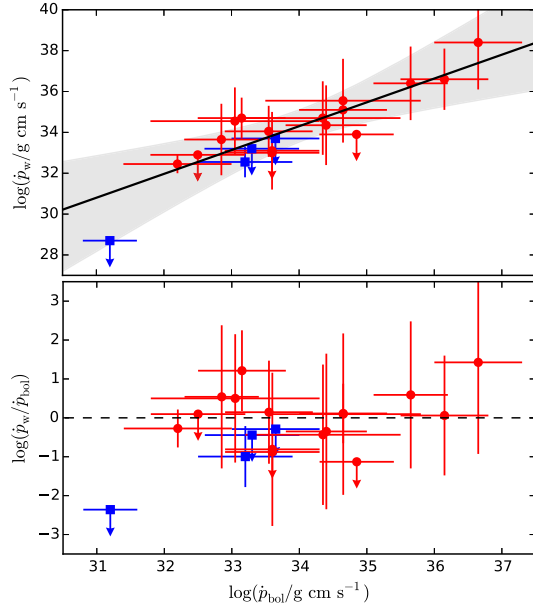


Figure 5. Comparison of the outflow (\dot{p}_w) and the radiation field (\dot{p}_{bol}) momentum rates. Top: a scatter plot showing the clear linear relationship between the two quantities (see Table 4 for parameters). Plot components are the same as previous. Bottom: constraint diagram showing the \dot{p}_w/\dot{p}_{bol} ratio for each source. Almost all of the AGNs in the sample are consistent with a ratio of unity (denoted by the dotted black line).

Another useful relationship can be obtained by multiplying equation (11) again by v_w :

$$L_w = b\tau \frac{v_w}{c} L_{Edd} = b \frac{\tau}{\lambda} \frac{v_w}{c} L_{bol}, \quad (12)$$

i.e. the wind kinetic power is predicted to be proportional to the bolometric luminosity of the AGN. In Fig. 3(d), we showed that L_w and L_{bol} are indeed correlated (Table 4), although the slope ($\beta \sim 1.5^{+1.0}_{-0.8}$) appears to be slightly steeper than (but still marginally consistent with) that predicted from equation (12). The steeper slope here is probably due to the $v_w \sim L_{bol}^{0.5}$ relationship observed before. Indeed, taking this into account yields $L_w \propto L_{bol}^{0.5} \times L_{bol} = L_{bol}^{1.5}$, which is exactly what we observe in these data.

Even so, it is worth noting that other acceleration mechanisms are not conclusively ruled out for these winds. If the winds are launched with zero initial radial velocity, then the terminal velocity of the wind in terms of the Eddington limit is $v_{inf} = \sqrt{2GM_{BH}/R_{launch}(f_\tau L_{bol}/L_{Edd} - 1)}$, where v_{inf} is the local terminal velocity and f_τ is the force multiplier ($\equiv 1$ for Thomson scattering). As a result, any wind accelerated purely by Thomson-scattering will not be able to greatly exceed the local terminal velocity without either (a) additional sources of opacity in the flow (such that $f_\tau > 1$), or (b) additional mechanisms transferring momenta to the flow. In this case, additional opacity in the flow may be attributed to the outflowing gas if it is stratified and contains clumps of lower ionization gas. This extra opacity could perhaps be associated with the partially covering gas required by many of our models in Paper I. Alternatively, magnetic processes may play a role in the initial acceleration which provides additional momentum to the flow, thereby allowing it to escape the system. Indeed, models have been proposed along these lines in the literature (e.g. Ohsuga et al. 2009; Ohsuga 2012; Ohsuga & Mineshige 2014). It is not possible to distinguish between these possibilities on the basis of existing data.

Table 4. Summary of linear regression parameters: $\log(y) = \alpha + \beta \log(x)$. Columns: (1) slope of the linear regression; (2) normalization/intercept of the best-fitting regression line; (3) standard deviation of the intrinsic scatter in the data; (4) best-fitting linear correlation coefficient. ± 1 denote perfect positive/negative correlations, respectively.

| Relation | | α | β | σ | R_p |
|-----------------|---------------------------|-------------------------|----------------------|---------------------|-------------------------|
| $\log(x)$ | $\log(y)$ | (1) | (2) | (3) | (4) |
| L_{bol} | N_H^a | $21.1^{+10.7}_{-10.5}$ | $0.0^{+0.3}_{-0.2}$ | $0.6^{+0.3}_{-0.2}$ | $0.11^{+0.59}_{-0.65}$ |
| | ξ^a | $0.3^{+7.1}_{-7.2}$ | $0.1^{+0.2}_{-0.2}$ | $0.4^{+0.6}_{-0.1}$ | $0.32^{+0.49}_{-0.60}$ |
| | v_w | $-13.6^{+12.4}_{-15.4}$ | $0.4^{+0.3}_{-0.2}$ | $0.6^{+0.3}_{-0.2}$ | $0.69^{+0.26}_{-0.47}$ |
| | r | $-8.4^{+38.3}_{-35.9}$ | $0.6^{+0.8}_{-0.7}$ | $0.8^{+0.8}_{-0.6}$ | $0.71^{+0.28}_{-0.93}$ |
| | \dot{M}_w | $-13.0^{+28.1}_{-35.4}$ | $0.9^{+0.8}_{-0.6}$ | $0.6^{+0.7}_{-0.4}$ | >0.34 |
| | L_w | $-23.5^{+23.6}_{-44.7}$ | $1.5^{+1.0}_{-0.8}$ | $1.0^{+1.2}_{-0.7}$ | >0.54 |
| | \dot{p}_w | $-18.1^{+28.8}_{-36.9}$ | $1.2^{+0.8}_{-0.7}$ | $0.7^{+0.9}_{-0.5}$ | >0.55 |
| \dot{p}_{bol} | r/r_s | $24.1^{+84.0}_{-34.3}$ | $-0.5^{+0.8}_{-0.8}$ | $0.8^{+0.9}_{-0.6}$ | $-0.60^{+0.99}_{-0.38}$ |
| | \dot{M}_w/\dot{M}_{Edd} | $0.4^{+29.5}_{-29.6}$ | $0.0^{+0.7}_{-0.7}$ | $0.5^{+0.7}_{-0.1}$ | $0.04^{+0.90}_{-0.99}$ |
| | L_w/L_{Edd} | $-27.0^{+34.0}_{-35.2}$ | $0.6^{+0.8}_{-0.8}$ | $0.8^{+0.9}_{-0.6}$ | $0.75^{+0.24}_{-1.02}$ |
| | \dot{p}_w/\dot{p}_{Edd} | $-12.4^{+30.5}_{-30.9}$ | $0.3^{+0.7}_{-0.7}$ | $0.6^{+0.8}_{-0.4}$ | $0.56^{+0.43}_{-1.32}$ |
| | \dot{p}_w | $-5.4^{+21.6}_{-27.9}$ | $1.2^{+0.8}_{-0.7}$ | $0.7^{+0.9}_{-0.5}$ | >0.55 |
| | | | | | |
| | | | | | |

Notes. All regressions were performed using the `linmix_err` Bayesian regression routine (Kelly 2007) which takes into account both measurement and upper limits in the y-variable. The noted best-fitting values correspond to the median value of the posterior probability distribution, as simulated from 100 000 MCMC realizations of the data, with errors taken as the range of values which encompass 90 per cent of the posterior probability. See Kelly (2007) for more details. Nominal values correspond to a fit to the entire sample, while those in brackets are for an analogous fitted with NGC 4395 removed. ^aRegression computed with lower limits excluded.

4.2.1 Estimating the wind opacity

Considering again the continuum-driven scenario, an interesting thing of note from Fig. 5(b) is that while all of the fast Fe K winds are consistent with the unity ratio within the errors, all of the slow winds fall below the limit (even when only the upper limits are considered). This possibly reflects subtle changes in the τ/λ ratio which allows for the possibility of continuum-driven winds in sub-Eddington AGN, whilst also allowing for subtle differences in the coupling efficiency of the gas between the two groups. As mentioned in Section 1, the AGNs considered here are relatively tightly clustered around a mean $\lambda \sim 0.1$ which means that we are unable to reliably test how the Eddington ratio affects the \dot{p}_w/\dot{p}_{bol} ratio. Nevertheless, we can still try to gauge how τ changes across the sample. From equation (11) we can see that when normalized to Eddington $\dot{p}_w/\dot{p}_{Edd} = 2b\tau$, with $b \sim \Omega/4\pi \approx 0.4$ (see Section 2.2) and therefore that $2b \simeq 1$ to within an order of magnitude. This means that $\dot{p}_w/\dot{p}_{Edd} \simeq \tau$, and that the \dot{p}_w/\dot{p}_{Edd} ratio can be in principle be used as a proxy for the gas opacity when the wind was launched. Fig. 6 shows how $\log(\dot{p}_w/\dot{p}_{Edd})$ varies with v_w . The higher velocity systems are largely consistent with $\dot{p}_w = \dot{p}_{Edd}$ ($\tau = 1$) within the errors, with a mean of $\langle \tau \rangle_{fast} \sim 0.3$, which is consistent with what is argued by King & Pounds (2003). In contrast, the slow systems (blue) all appear to have \dot{p}_w/\dot{p}_{Edd} ratios around an order of magnitude lower. Taken at face value, this possibly indicates that they have a lower opacity to scattering. There also appears to be a transitional phase between the two groups, with $v_w \sim$ a few thousand km s^{-1} , which may be consistent to the idea of stratified winds launched over a wide range of radii with varying velocities (see Tombesi et al. 2013). It is worth noting, however, that the opacity

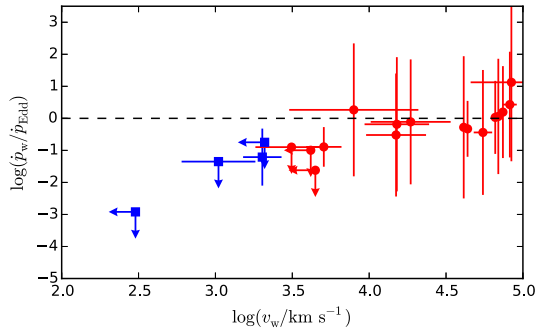


Figure 6. Plot showing $\log(\dot{p}_w/\dot{p}_{\text{Edd}})$ versus $\log(v_w/\text{km s}^{-1})$ which can be used as a diagnostic for τ (see text for details). The majority of fast AGNs ($v_w > 3000 \text{ km s}^{-1}$; red points) are consistent with $\dot{p}_w = \dot{p}_{\text{Edd}}$, thereby implying that $\tau = 1$, while all of the slow winds (blue points) are well below a ratio of unity.

estimated from $\dot{p}_w/\dot{p}_{\text{Edd}}$ could be lower than that of the bulk flow if (a) we view the wind from an acute angle where the observed velocity is not representative of the bulk flow velocity, and (b) the wind has a component of completely ionized gas which contributes to the scattering opacity but has no spectroscopic signatures. One possibility is that all of the winds have $\tau \simeq 1$ when integrated over all lines of sight and across all ionization phases. Alternatively, the fact that the slow winds have $\tau \ll 1$, coupled with the fact that they are preferentially observed in lower luminosity AGN, may indicate that they are accelerated by another mechanism such as magnetic pressure, for example, which may have a larger contribution at lower luminosities (e.g. see Ohsuga et al. 2009; Ohsuga 2012).

Regardless, the fast winds – which are the ones most likely to affect their host galaxy due to their large velocity and therefore the most interesting in terms of AGN–host-galaxy feedback – are consistent with a wind launched by continuum scattering at the Eddington limit (e.g. King & Pounds 2003).

4.3 Energetic significance and feedback implications

The independent *Suzaku* and *XMM–Newton* wind samples have conclusively shown highly ionized Fe K winds are both frequently observed in AGN and that they contain a large amount of mechanical energy. However, a fundamental question still remains: Is the energy imparted by the mass flow likely to play a role in terms of AGN–host-galaxy feedback scenarios? AGN feedback models have postulated that ~ 5 per cent of an AGN’s bolometric radiative output needs to be converted to mechanical power in order to have a notable effect on the growth of a central SMBH and the host galaxy (e.g. Di Matteo et al. 2005). However, recent numerical simulations by Hopkins & Elvis (2010) have shown that should the effects of ‘secondary’ feedback such as cloud ablation be taken into account, the required energy can be significantly lower, of the order of ~ 0.5 per cent. Moreover, hydrodynamic simulations have also shown that even relatively modest Fe K outflows (e.g. those with only $v_w \sim 0.01c$, $\dot{M}_w \sim 0.01 M_\odot \text{ yr}^{-1}$ and $L_w \sim 10^{44} \text{ erg s}^{-1}$; see Wagner, Umemura & Bicknell 2013) can impart significant feedback upon the host galaxy on $\sim \text{kpc}$ scales should these effects be taken into account. Investigating what fraction of an AGN’s bolometric output is conveyed through mechanical power therefore provides a means of qualitatively assessing the likely energetic significance of the mass flow. A plot of $L_{\text{bol}} - (L_w/L_{\text{bol}})$ to this effect is shown in Fig. 7.

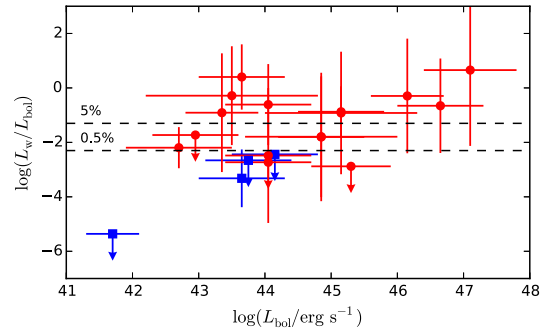


Figure 7. Diagnostic plot comparing the ratio of L_w/L_{bol} to the L_{bol} of each AGN in the *Suzaku* sample. The horizontal lines denote L_w/L_{bol} fractions of 0.5 per cent and 5 per cent which are the theoretical thresholds for feedback. Almost all of the fast absorption systems (red points) have L_w in excess of 0.5 per cent of L_{bol} which suggests that they may be significant in terms of AGN–host-galaxy feedback scenarios.

A large fraction of the outflows have a kinetic luminosity either in excess of or consistent with the 0.5 per cent of L_{bol} , with 10/20 also exceeding 5 per cent as well. The mean kinetic power considering only the fast winds is $\langle L_w/L_{\text{bol}} \rangle \approx 7$ per cent. This is well above the supposed ~ 0.5 per cent threshold for significant feedback. In the most conservative case, i.e. by considering only L_w^{min} given the available errors, we estimate $\langle L_w/L_{\text{bol}} \rangle \approx 0.2$ per cent. This is comparable to the ~ 0.5 per cent threshold and consistent with what was found by Tombesi et al. (2012a) on the basis of the *XMM–Newton* outflow sample. Therefore, even in the most conservative case, the energetic output of the fast *Suzaku*-detected Fe K winds is comparable to the threshold thought necessary for feedback. Importantly, these estimates are derived solely from the wind kinematics and are therefore independent of the initial driving mechanism.

4.4 Relation to $M_{\text{BH}} - \sigma_*$ in quiescent galaxies

The FeK winds detected here may therefore impart enough energy into the host galaxy to play a role in galaxy-scale feedback. With this in mind, it is useful to highlight the recent work of McQuillin & McLaughlin (2013). To set the context of their work, it is first useful to explain the difference between momentum- and energy-driven outflows. As a wind propagates into the host galaxy it will sweep up a shell of ambient gas, the dynamics of which are determined by how efficiently the shocked wind material behind the shell can cool. Initially, the shocked gas can cool efficiently (by Compton scattering) so the shell stays thin and is driven outwards directly by the force of the wind (cf. our equation 11; also King & Pounds 2003; King 2010; McQuillin & McLaughlin 2012). This is momentum-driven feedback. As the outflow reaches larger distances, however, the cooling efficiency drops and the shocked gas becomes geometrically thick and hot. The outward force on the shell then comes from the wind thrust in equation (11) mediated by the thermal pressure and work done on and by the shocked gas. This is energy-driven feedback (e.g. King 2003, 2005; Zubovas & King 2012; McQuillin & McLaughlin 2013). The transition between the two phases can occur at relatively small radii (see McQuillin & McLaughlin 2013), and energy-driven outflows are therefore expected to be dominant on large galactic scales.

McQuillin & McLaughlin (2013) consider purely energy-driven outflows and investigate their link to the observed $M_{\text{BH}} - \sigma_*$ relationship. They showed that, for an energy-driven shell to reach the

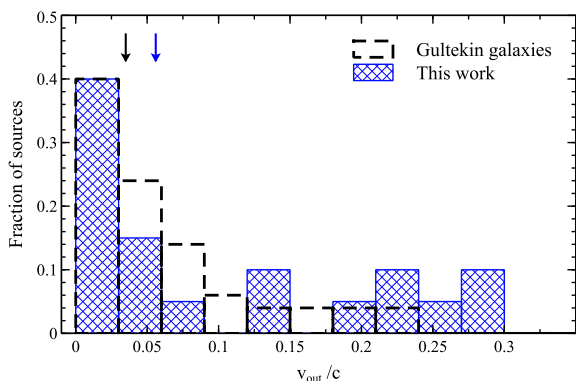


Figure 8. Histogram comparing the *Suzaku*-observed distribution of outflow velocities (blue cross-hatched area) to those inferred from local quiescent galaxies by McQuillan & McLaughlin (2013) who assume that their observed $M_{\text{BH}}-\sigma_*$ relation (cf. Gültekin et al. 2009) is induced by energy-driven feedback, with $M_{\text{BH}}v_w \propto \sigma_*^5$. The blue and black arrows correspond to the median outflow velocities of $v_w \sim 0.056c$ and $0.035c$ in the *Suzaku*-observed AGN and quiescent sub-samples, respectively. The two samples are formally indistinguishable by the KS test. See text for further discussion.

escape speed from an isothermal sphere of dark matter and gas with velocity dispersion σ requires a critical BH mass given by

$$\left(\frac{v_w}{c}\right) \left(\frac{M_{\text{BH}}}{10^8 M_\odot}\right) = 6.68 \times 10^{-2} \left(\frac{\sigma_*}{200 \text{ km s}^{-1}}\right)^5, \quad (13)$$

such that the velocity of the wind from the BH (v_w) enters explicitly. McQuillan & McLaughlin (2013) used this equation to infer the wind speed that would have been needed during the active protogalactic phase of 51 quiescent ($z \approx 0$) galaxy spheroids (taken from the sample of Gültekin et al. 2009) in order to explain the observed values of M_{BH} and σ in these galaxies as a result of (energy-driven) feedback. Their main result is a distribution of BH wind speeds, derived in essence from the scatter of the $M_{\text{BH}}-\sigma_*$ relation, that bears a remarkable resemblance to the empirical distributions of v_w in the AGN samples of both Tombesi et al. (2010a) and Paper I.

Fig. 8 compares the velocity distribution inferred by McQuillan & McLaughlin (2013) to the one that we measured for *Suzaku*-detected Fe K outflows (note that we consider all of the outflows in this panel). As noted by McQuillan & McLaughlin (2013), there is a close resemblance between the two, with both samples tracing a similar overall distribution. Their overall range are similar, as is their median: $v_w \sim 0.035c$ in the quiescent galaxies, and $\sim 0.056c$ for the *Suzaku*-detected outflows (see Paper I). Formally, the two distributions cannot be distinguished: a KS test only rules their being from the same parent population at $P_{\text{KS}} \sim 75$ per cent confidence. The fact that there are similarities between the inferred and observed v_w distributions at all is in itself extremely interesting, not only because the independent samples of quiescent galaxies and AGNs are drawn from two different sub-populations of galaxy, but because it implies a close relationship between the emergent wind and the evolution of the host galaxy. Overall, the work of McQuillan & McLaughlin (2013) is encouraging for the general notion that high-velocity outflows contribute to the feedback which is invoked to explain the $M_{\text{BH}}-\sigma_*$ relationship in local quiescent galaxies, and, more importantly, for the idea that AGN play a key role in galaxy formation and evolution.

5 SUMMARY AND CONCLUSIONS

In Paper I, we formed a detailed spectroscopic study of Fe K absorption in a heterogeneous sample of 52 *Suzaku*-observed AGN. There, we found that ~ 40 per cent of the sources in the sample harboured statistically significant Fe XXVI He α and/or Fe XXVI Ly α absorption lines with velocities ranging from a few thousand km s^{-1} to $\sim 0.3c$. We also measured their column density ionization parameter of the absorber using the *XSTAR* photoionization code. Here, we build upon these results to compute the location r , mass outflow rate \dot{M}_w , kinetic power L_w and momentum flux \dot{p}_w of the implied wind, and use these results to inform an exploratory analysis of how the wind parameters scale with the bolometric luminosity of their host AGN. We focus our attention on the fastest winds, defined here as $v_w > 0.01c$, because these ones likely represent genuine disc-winds which are likely distinct from the slower (and more distant) warm absorber. The main results of our study are summarized below.

(i) The mean radial distance to the fast winds is $\langle r \rangle \sim 10^{15-17}$ cm (typically $\sim 10^{2-4} r_s$) from the BH. This corresponds to ~ 0.0003 – 0.03 pc, such that the Fe K absorbers are located much closer to the BH than the traditional parsec-scale warm absorber. The mean wind mass outflow rate and mean kinetic power are constrained to $\langle \dot{M}_w \rangle \sim 10^{24-26} \text{ g s}^{-1}$ (~ 0.01 – $1 \dot{M}_w \text{ yr}^{-1}$) and $\langle L_w \rangle \sim 10^{43-44} \text{ erg s}^{-1}$, respectively. The average upper limit is $\sim \dot{M}_{\text{Edd}}$, but is still an appreciable ~ 1 per cent of \dot{M}_{Edd} if only the lower limits are considered, whereas the kinetic power is constrained to $\sim (0.01$ – $0.1)L_{\text{Edd}}$. These properties are consistent with wind at a substantial fraction of the Eddington limit which carries a large amount of mass into the host galaxy.

(ii) The Fe K absorber column density and ionization parameter are largely independent of L_{bol} but there is a significant correlation between the wind velocity v_w and L_{bol} , with a slope $\beta = 0.4^{+0.3}_{-0.2}$. This tentatively indicates that more luminous AGN launch winds with a larger observed velocity along the LOS.

(iii) The winds in more luminous AGN contain more material, and they are subsequently more energetic. The median slope of the $L_{\text{bol}}-L_w$ and $L_{\text{bol}}-\dot{p}_w$ regressions are formally consistent with unity.

(iv) There are no significant correlations in the $L_{\text{bol}}-(r/r_s)$ and $L_{\text{bol}}-(\dot{M}_w/\dot{M}_{\text{Edd}})$ planes. Therefore, the winds are located at similar distances in Schwarzschild units and contain a similar fraction of \dot{M}_{Edd} .

(v) The momentum flux of the radiation field (\dot{p}_{bol}) and that of the wind (\dot{p}_w) are strongly correlated, with a slope $\beta = 1.2^{+0.8}_{-0.7}$. This is quantitatively consistent with what is expected from a continuum-driven wind.

(vi) A significant fraction of the sample (17/20, 85 per cent) exceeded the minimum $L_w/L_{\text{bol}} \sim 0.5$ per cent threshold thought necessary for feedback (Hopkins & Elvis 2010), while 9/20 (45 per cent) also exceed the less conservative ~ 5 per cent threshold as well (Di Matteo et al. 2005). In the most conservative case, the mean L_w/L_{bol} ratio is ~ 0.2 per cent. This suggests that the winds may be sufficiently energetic in terms of feedback.

These results enforce those recently obtained with *XMM-Newton* and provide additional evidence in favour of Fe K absorption being a prevalent feature of the AGN X-ray spectrum, and that the ensuing wind may be energetically significant in terms of feedback.

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